

Resonant non-Gaussianity with equilateral properties

Rhiannon Gwyn[#], Markus Rummel^{*} and Alexander Westphal[†]

[#]*AEI Max-Planck-Institut für Gravitationsphysik, D-14476 Potsdam, Germany*

^{*}*II. Institut für Theoretische Physik der Universität Hamburg, D-22761 Hamburg, Germany*

[†]*Deutsches Elektronen-Synchrotron DESY, Theory Group, D-22603 Hamburg, Germany*

We discuss the effect of superimposing multiple sources of resonant non-Gaussianity, which arise for instance in models of axion inflation. The resulting sum of oscillating shape contributions can be used to “Fourier synthesize” different non-oscillating shapes in the bispectrum. As an example we reproduce an approximately equilateral shape from the superposition of $\mathcal{O}(10)$ oscillatory contributions with resonant shape. This implies a possible degeneracy between the equilateral-type non-Gaussianity typical of models with non-canonical kinetic terms, such as DBI inflation, and an equilateral-type shape arising from a superposition of resonant-type contributions in theories with canonical kinetic terms. The absence of oscillations in the 2-point function together with the structure of the resonant N -point functions, imply that detection of equilateral non-Gaussianity at a level greater than the PLANCK sensitivity of $f_{NL} \sim \mathcal{O}(5)$ will rule out a resonant origin. We comment on the questions arising from possible embeddings of this idea in a string theory setting.

November 2, 2012

Contents

1	Introduction	2
2	A sum of resonant bispectra	4
2.1	Potential, solution and slow-roll parameters	4
2.2	The power spectrum	6
2.3	The bispectrum	8
3	Equilateral features from summed resonant non-Gaussianity	10
3.1	Equilateral and local shapes for $x_- \rightarrow 0$	12
3.2	Scale invariance vs. scale dependence	13
3.3	Fourier analysis	15
3.4	Constraints from the power spectrum	18
4	Possible avenues for embedding: Axion monodromy inflation	19
4.1	Axion monodromy inflation	19
4.2	Periodic corrections to the potential	21
5	Conclusions	23
A	Computing the mode functions and bispectrum	25
A.1	The mode functions	25
A.2	Mode functions in the bispectrum	28
A.3	Squeezed limit consistency relation	28

1 Introduction

Recent years have seen the advent of precise observations of the cosmic microwave background radiation (CMB) [1, 2, 3], and the redshift-distance relation for large samples of distant type IA supernovae [4, 5], as well as a host of additional increasingly precise measurements such as baryon acoustic oscillations [6] or the determination of the Hubble parameter H_0 by the Hubble Space Telescope key project [7]. These results are so far in concordance with a Universe very close to being spatially flat, as well as with a pattern of coherent acoustic oscillations in the early dense plasma which was seeded by an almost scale-invariant spectrum of super-horizon size curvature perturbations with Gaussian distribution. This structure is a direct consequence of a wide class of models of cosmological inflation driven by the potential energy of a single scalar field with canonically normalized kinetic energy in (strict) slow-roll. In these models of inflation the spectrum of scalar curvature perturbations generated during the quasi-de-Sitter phase is necessarily Gaussian, with 3-point interactions generating very small non-Gaussian deviations $f_{NL} = \mathcal{O}(\epsilon, \eta)$ of order of the slow-roll parameters during inflation [8]. Consequently there are large classes of inflationary models which generate sizable levels of non-Gaussianity and are characterized by departures from one or more of the defining properties of canonically normalized single-field slow-roll inflation.

Non-Gaussianity can arise from a non-canonical kinetic term for a single scalar field characterized by the presence of higher-derivative terms, from controlled non-permanent violations of slow-roll such that the overall inflationary behaviour is preserved, or from the presence of several light scalar fields during inflation. A series of results have established that these different mechanisms of generating a large 3-point function and its associated non-Gaussianity lead to different ‘shapes’, the distributions of the amount of non-Gaussianity over the momentum-conservation dictated triangular domain of the three fluctuation momenta in the 3-point function. These shape functions allow for discrimination among many different classes of inflationary models, provided one has sufficient measurement accuracy for the shape, and one has shown the existence of a complete orthonormal system of shapes of non-Gaussianity in the space of inflation models. For two recent reviews see e.g. [9, 10].

In this article, we explore the degeneracy, at the level of non-Gaussianity, of different classes of inflationary models. Specifically we study degeneracy between models of inflation characterized by so-called non-canonical kinetic terms such as DBI inflation [11, 12] and

theories which are purely canonical. By non-canonical kinetic terms we mean terms of the form X^n , where $X = \frac{1}{2}\partial^\mu\phi\partial_\mu\phi$, so that a non-canonical Lagrangian is given by some general $P(X, \phi)$ [13]. These are thus a subset of the theories described by the Horndeski action, which is the most general action for a single field with at most two time derivatives acting on it at the level of the equation of motion, i.e. the most general action for a single scalar field without ghosts or Ostrogradski instabilities (See e.g. [14] and [15]).

Theories in this class give rise to an equilateral-type non-Gaussianity inversely proportional to the reduced speed of sound, c_s (see e.g. [16]):

$$f_{NL}^{equil} \sim \frac{1}{c_s^2}, \quad (1.1)$$

where

$$c_s^2 = \left(1 + 2X \frac{P_{XX}}{P_X}\right)^{-1}. \quad (1.2)$$

By contrast, canonical slow-roll single scalar field models give rise to Gaussian spectra (with corrections of order the slow-roll parameters) for which all odd-point correlation functions vanish and all even-point correlations are given in terms of the two-point functions. How then could there be any observational degeneracy between a model of non-canonical inflation and a canonical single scalar field model of inflation? In fact, single scalar field models with canonical kinetic term give rise to (approximately) Gaussian spectra only as long as the slow-roll conditions are always satisfied. As argued in [17], these theories can generate observably large non-Gaussianities when there are features in the potential. The two classes of such models considered in [17] are: 1) potentials with steps resulting in a localized violation of slow-roll, and 2) potentials with a small modulation which imply a large variation $\dot{\eta}$ for small ϵ and η . Because one of the terms in the three-point function is proportional to $\epsilon\dot{\eta}$, this can give rise to an appreciable non-Gaussianity with a specific shape, termed resonant non-Gaussianity [18], which is orthogonal to the local and equilateral shapes. Resonant non-Gaussianity comes with the additional property of allowing for very efficient calculability of the N -point functions up to $N = 10 \dots 20$ [19].

However, as we shall show, a closely related model, in which a sum of such modulations in the potential is present, leads to a very different non-Gaussianity. Its shape is of the equilateral type in the sense that its momentum average has a large overlap with the standard equilateral one, despite the lack of non-canonical kinetic terms. Yet, periodic scale dependent

features remain. Thus observation of equilateral non-Gaussianity might not necessarily imply that non-canonical kinetic terms were present or relevant in the underlying model of inflation. This degeneracy at the level of the three-point function leads us to conclude that it may not be as easy as previously hoped to distinguish amongst single-field models of inflation. However, we then clarify that the non-observation of oscillating contributions to the 2-point function, and the structure of the resonant N -point functions place tight constraints on the maximum amount of resonantly generated non-Gaussianity with equilateral characteristics. In particular, this implies that detection of equilateral non-Gaussianity at a level greater than the PLANCK sensitivity of $f_{NL} \sim \mathcal{O}(5)$ will rule out a resonant origin [20]. In case a detection of non-Gaussianity eludes PLANCK, future 21cm observations may provide a detection capability of $f_{NL} \gtrsim \mathcal{O}(0.01)$ down to the slow-roll level itself due to the extremely large number of modes ($\mathcal{O}(10^{16})$) available [21, 22]. For equilateral-type non-Gaussianities of this magnitude there is then a degeneracy between a possible non-canonical and a resonant origin.

2 A sum of resonant bispectra

2.1 Potential, solution and slow-roll parameters

From an effective field theory point of view, we can begin with a model of a single scalar field with canonical kinetic term and modulated potential

$$V(\phi) = V_0(\phi) + \sum_i A_i \cos\left(\frac{\phi + c_i}{f_i}\right), \quad (2.3)$$

where we have generalized the modulated potential in [18] to the case where the modulation is a series of terms with varying phases. For suitable coefficients and values of f_i , the sum remains a small perturbation on the potential $V_0(\phi)$ (see Figure 1). We will show that slow-roll inflation for this theory is supported for sufficiently large f_i , while a large variation in the slow-roll parameters is possible. In Sections 2.2 and 2.3 we give the power spectrum and bispectrum respectively, finding that the bispectrum is given by a series of the resonant bispectra found for a singly modulated potential in [18]. We shall follow closely the analysis in [18]. The details of our calculations are given in Appendix A.

The equation of motion for the inflaton ϕ with potential eq. (2.3) is

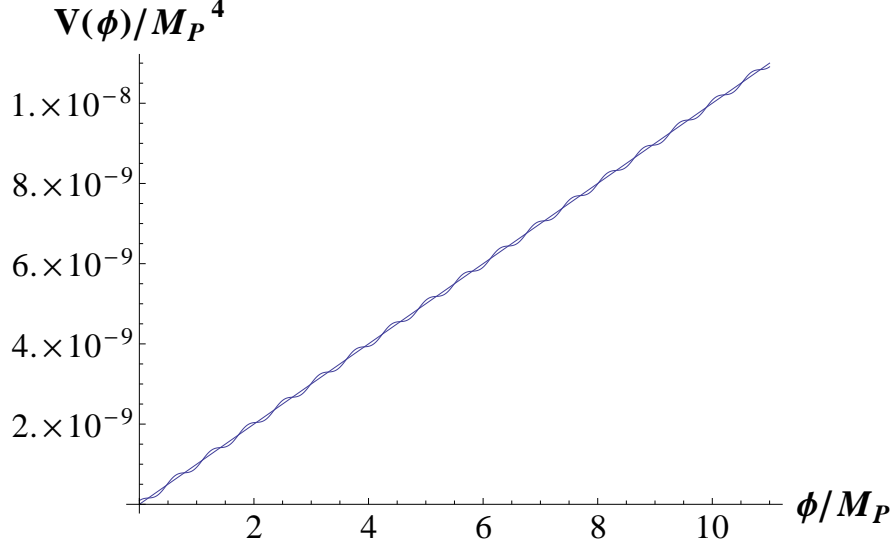


Figure 1: Modulated linear potential, for a convenient choice of c_i, f_i values.

$$\ddot{\phi} + 3H\dot{\phi} + V'_0(\phi) = \sum_i \frac{A_i}{f_i} \sin\left(\frac{\phi + c_i}{f_i}\right). \quad (2.4)$$

As in [18] we treat the oscillations as a perturbation and write $\phi = \phi_0 + \phi_1 + \dots$ where ϕ_0 is the solution in the absence of modulations, given by

$$\dot{\phi}_0 = -\frac{V'_0(\phi_0)}{\sqrt{3V_0(\phi_0)}}. \quad (2.5)$$

The equation of motion for ϕ_1 , written in terms of derivatives with respect to ϕ_0 (denoted by a prime) is

$$\phi_1'' - \frac{3}{\sqrt{2\epsilon_{V_0}}} \phi_1' - \left(\frac{3}{2} - \frac{3\eta_{V_0}}{2\epsilon_{V_0}}\right) \phi_1 = \frac{3}{2\epsilon_{V_0}V_0} \sum_i \frac{A_i}{f_i} \sin\left(\frac{\phi_0 + c_i}{f_i}\right), \quad (2.6)$$

where $V_0 = V_0(\phi_0)$, and ϵ_{V_0} and η_{V_0} are the canonical slow-roll parameters evaluated for the potential V_0 . Note that $\epsilon'_{V_0} = \sqrt{2\epsilon_{V_0}}(\eta_{V_0} - 2\epsilon_{V_0})$. The solution is a generalization of that found in [18]; to first order in slow-roll parameters and assuming $f_i \ll \sqrt{2\epsilon_\star} \forall i$, it is given by

$$\phi_1(t) = -\frac{3}{2\epsilon_\star V_0(\phi_\star)} \sum_i A_i f_i \sin\left(\frac{\phi_0 + c_i}{f_i}\right), \quad (2.7)$$

where the subscript \star indicates that $\epsilon = \epsilon_{V_0}(\phi_\star)$ is evaluated at horizon exit. This solution works because $\phi_1 = \sum \phi_i$, where each ϕ_i satisfies the EOM for a specific f_i . $\phi_0(t)$ is time

dependent but, for the period of interest, in which modes observed today exit the horizon, we can approximate it as being given by the value at horizon crossing, i.e. ϕ_* , except in the oscillating term. In other words we have assumed that the dominant variation in ϕ_1 is due to the modulating terms. Defining

$$b_i^* \equiv \frac{A_i}{V_0'(\phi_*)f_i}, \quad (2.8)$$

we have that the full solution is well approximated by

$$\phi(t) = \phi_0(t) - \sum_i \frac{3b_i^* f_i^2}{\sqrt{2\epsilon_*}} \sin\left(\frac{\phi_0 + c_i}{f_i}\right). \quad (2.9)$$

The slow-roll parameters $\epsilon = -\frac{\dot{H}}{H^2}$ and $\delta = \frac{\ddot{H}}{2H\dot{H}}$ can be calculated for the potential eq. (2.3) and are found to be:

$$\begin{aligned} \epsilon &= \epsilon_* + \left(\frac{V_0'}{V_0}\right)^2 \phi_1', \\ \epsilon &= \epsilon_* - 3\sqrt{2\epsilon_*} \sum_i b_i^* f_i \cos\left(\frac{\phi_0 + c_i}{f_i}\right) + \mathcal{O}((b_i^*)^2), \\ \epsilon &= \epsilon_0 + \epsilon_1 + \mathcal{O}((b_i^*)^2); \\ \delta &= \epsilon_* - \eta_* - \sqrt{2\epsilon_*} \phi_1'' + \epsilon_*(\eta_* - \epsilon_*)\phi_1, \\ \delta &= \epsilon_* - \eta_* - \sum_i 3b_i^* \sin\left(\frac{\phi_0 + c_i}{f_i}\right) + \mathcal{O}((b_i^*)^2), \\ \delta &= \delta_0 + \delta_1 + \mathcal{O}((b_i^*)^2). \end{aligned} \quad (2.10)$$

The SR parameter of interest for observation is $\frac{\dot{\delta}_1}{H}$ where $\delta_0 = \epsilon_* - \eta_*$ is the value of δ coming from the unmodulated potential V_0 . A large $\frac{\dot{\delta}}{H}$ is what is responsible for a large resonant non-Gaussianity in [18]. Here, we find

$$\frac{\dot{\delta}_1}{H} = \sum_i \frac{\sqrt{2\epsilon_*}}{f_i} 3b_i^* \cos\left(\frac{\phi_0 + c_i}{f_i}\right). \quad (2.11)$$

Note that $\frac{\sqrt{2\epsilon_*}}{f_i}$ is large for all i for the solution eq. (2.9).

2.2 The power spectrum

As in [18], the Mukhanov-Sasaki equation for the mode function \mathcal{R}_k of the curvature perturbation in our case is given by (see the Appendix for details)

$$\frac{d^2 \mathcal{R}_k}{dx^2} - \frac{2(1 + \delta_1(x))}{x} \frac{d\mathcal{R}_k}{dx} + \mathcal{R}_k = 0, \quad (2.12)$$

where $x = -k\tau$ for τ the conformal time. We have neglected ϵ_0 and δ_0 and are working only to leading order in ϵ_\star . The form of this equation is the same as in [18], but δ_1 is now defined in eq. (2.10) and given by

$$\delta_1 = -3 \sum_i b_i^\star \sin\left(\frac{\phi_0 + c_i}{f_i}\right). \quad (2.13)$$

We will find a sum of integrals each encountering resonance at a value $x_{\text{res},i}$. As in the case of a single modulation, the effect of δ_1 is negligible both at early times $x \gg x_{\text{res},i}$ and at late times $x \ll x_{\text{res},i}$. We therefore look for a solution of the form

$$\mathcal{R}_k(x) = \mathcal{R}_{k,0}^{(o)} \left[i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(1)}(x) - c_k^{(-)}(x) i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(2)}(x) \right], \quad (2.14)$$

where $\mathcal{R}_{k,0}^{(o)}$ is the value of $\mathcal{R}_k(x)$ outside the horizon in the absence of modulations, $H_{3/2}^{(1,2)}(x)$ are Hankel functions and $c_k^{(-)}(x)$ vanishes at early times. Then the Mukhanov-Sasaki equation gives an equation governing the time evolution of $c_k^{(-)}(x)$ which is given by (for $f_i \ll \sqrt{2\epsilon_\star}$ for all i)

$$\frac{d}{dx} \left[e^{-2ix} \frac{d}{dx} c_k^{(-)}(x) \right] = -2i \frac{\delta_1(x)}{x}. \quad (2.15)$$

To solve this we need δ_1 as a function of the conformal time τ . In terms of τ the background equation of motion is given by

$$\frac{d\phi_0}{d\ln(-\tau)} = \sqrt{2\epsilon_\star}, \quad (2.16)$$

with solution $\phi_0(\tau) = \phi_\star + \sqrt{2\epsilon_\star} \ln \frac{\tau}{\tau_\star}$. This can be written as $\phi_0(x) = \phi_k + \sqrt{2\epsilon_\star} \ln x$, where $\phi_k = \phi_\star - \sqrt{2\epsilon_\star} \ln \frac{k}{k_\star}$ is the value of the field when the mode with comoving momentum k exits the horizon. Then, writing $c_k^{(-)}(x) = \sum_i c_{k,i}^{(-)}(x)$, we need to solve

$$\frac{d}{dx} c_{k,i}^{(-)}(x) = -\frac{6ib_i^\star f_i}{\sqrt{2\epsilon_\star}} e^{2ix} \cos\left(\frac{\phi_k}{f_i} + \frac{c_i}{f_i} + \frac{\sqrt{2\epsilon_\star}}{f_i} \ln x\right). \quad (2.17)$$

Thus each $c_{k,i}^{(-)}(x)$ behaves like $c_k^{(-)}(x)$ in [18] but with $b^\star = b_i^\star$, $f = f_i$ and resonance at different x values $x_{\text{res},i} = (2f_i\phi_\star)^{-1}$.

For each term we can use the stationary phase approximation to find

$$c_{k,i}^{(-)} = 3b_i^\star \sqrt{\frac{\pi}{2}} \left(\frac{f_i}{\sqrt{2\epsilon_\star}} \right)^{1/2} e^{-i(\frac{\phi_k + c_i}{f_i})}, \quad (2.18)$$

$$\Rightarrow \mathcal{R}_k(x) = \mathcal{R}_{k,0}^{(o)} \left[i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(1)}(x) - \sum_i \frac{3ib_i^\star \pi}{2} \sqrt{\frac{f_i}{\sqrt{2\epsilon_\star}}} e^{-i(\frac{\phi_k + c_i}{f_i})} x^{3/2} H_{3/2}^{(2)}(x) \right]. \quad (2.19)$$

Outside the horizon, $k/aH \ll 1$ i.e. $x = -k\tau = k/aH \ll 1$ (using the definition of conformal time and the fact that during inflation H is constant). Then we find

$$\mathcal{R}_k(x) = R_{k,0}^{(0)} \left[1 + \sum_i 3b_i^* \sqrt{\frac{f_i \pi}{2\sqrt{2}\epsilon_\star}} \cos\left(\frac{\phi_k + c_i}{f_i}\right) - i \sum_i 3b_i^* \sqrt{\frac{f_i \pi}{2\sqrt{2}\epsilon_\star}} \sin\left(\frac{\phi_k + c_i}{f_i}\right) + \mathcal{O}(x^3) \right], \quad (2.20)$$

and

$$|\mathcal{R}_k^{(0)}|^2 = |\mathcal{R}_{k,0}^{(0)}|^2 [1 + \delta n_s], \quad (2.21)$$

with

$$\delta n_s = \sum_i 3b_i^* \left(\frac{2\pi f_i}{\sqrt{2}\epsilon_\star} \right)^{1/2} \cos\left(\frac{\phi_k + c_i}{f_i}\right), \quad (2.22)$$

which is just the generalization of (2.30) in [18] one might have expected.

2.3 The bispectrum

Next we ask if this generalization carries through to the bispectrum as calculated in [18]. As in the case of a single modulation, the leading contribution comes from

$$H_I(t) \supset - \int d^3x a^3(t) \epsilon(t) \dot{\delta}(t) \mathcal{R}^2(\mathbf{x}, t) \dot{\mathcal{R}}(\mathbf{x}, t), \quad (2.23)$$

where to linear order in b_i^* we can use the approximations $\epsilon \approx \epsilon_\star, \delta \approx \delta_1$ and use the unperturbed mode functions¹

$$\mathcal{R}_k(x) = \mathcal{R}_{k,0}^{(o)} i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(1)}(x). \quad (2.24)$$

The analysis of [18] therefore carries over to our case, with the three-point function given by

$$\langle \mathcal{R}(\mathbf{k}_1, t) \mathcal{R}(\mathbf{k}_2, t) \mathcal{R}(\mathbf{k}_3, t) \rangle = (2\pi)^7 \Delta_{\mathcal{R}}^4 \frac{1}{k_1^2 k_2^2 k_3^2} \delta^3(\mathbf{k}_1 + \mathbf{k}_2 + \mathbf{k}_3) \frac{\mathcal{G}(k_1, k_2, k_3)}{k_1 k_2 k_3}. \quad (2.25)$$

for

$$\frac{\mathcal{G}(k_1, k_2, k_3)}{k_1 k_2 k_3} = \frac{1}{8} \int_0^\infty dX \frac{\dot{\delta}_1}{H} e^{-iX} \left[-i - \frac{1}{X} \sum_{i \neq j} \frac{k_i}{k_j} + \frac{i}{X^2} \frac{K(k_1^2 + k_2^2 + k_3^2)}{k_1 k_2 k_3} \right] + c.c., \quad (2.26)$$

¹To see why we can use the unperturbed mode functions, see Appendix A.2.

where $K \equiv k_1 + k_2 + k_3$, $X \equiv -K\tau$, and $\dot{\delta}_1$ is given by eq. (2.11). Equivalently this can be written

$$\frac{\mathcal{G}(k_1, k_2, k_3)}{k_1 k_2 k_3} = \frac{1}{8} \sum_i \int_0^\infty dX \frac{\dot{\delta}_{1,i}}{H} e^{-iX} \left[-i - \frac{1}{X} \sum_{i \neq j} \frac{k_i}{k_j} + \frac{i}{X^2} \frac{K(k_1^2 + k_2^2 + k_3^2)}{k_1 k_2 k_3} \right] + c.c. \quad (2.27)$$

where

$$\frac{\dot{\delta}_{1,i}}{H} = \frac{\sqrt{2\epsilon_\star}}{f_i} 3b_i^\star \cos\left(\frac{\phi_0 + c_i}{f_i}\right) = \frac{\sqrt{2\epsilon_\star}}{f_i} 3b_i^\star \cos\left(\frac{\phi_K + \sqrt{2\epsilon_\star} \ln X + c_i}{f_i}\right). \quad (2.28)$$

Now, for each integral in the sum the biggest contribution comes from $X_{\text{res},i} = \frac{\sqrt{2\epsilon_\star}}{f_i}$. This can be seen by defining

$$\begin{aligned} \mathcal{I}_{K,i} &= \frac{3b_i^\star \sqrt{\epsilon_\star}}{f_i} \int_0^\infty dX e^{-iX} \cos\left(\frac{\phi_K + \sqrt{2\epsilon_\star} \ln X + c_i}{f_i}\right) \\ &= \frac{3b_i^\star \sqrt{2\epsilon_\star}}{2f_i} \int_0^\infty dX \left(e^{i(\phi_K + c_i)/f_i} e^{-i(X - \frac{\sqrt{2\epsilon_\star} \ln X}{f_i})} + e^{-i(\phi_K + c_i)/f_i} e^{-i(X + \frac{\sqrt{2\epsilon_\star} \ln X}{f_i})} \right). \end{aligned} \quad (2.29)$$

The phase of the first term can be zero while the phase of the second cannot, which means that the first term will dominate, with its dominant contribution at $X_{\text{res},i} = \frac{\sqrt{2\epsilon_\star}}{f_i}$. This confirms that the main contribution to the bispectrum arises when the modes are still deep in the horizon. In this regime, we can approximate $1/X$ and $1/X^2$ by $1/X_{\text{res},i}$ and $1/X_{\text{res},i}^2$ in eq. (2.27) and we have

$$\frac{\mathcal{G}(k_1, k_2, k_3)}{k_1 k_2 k_3} = \frac{1}{4} \sum_i \left[\text{Im} \mathcal{I}_{K,i} - \frac{1}{X_{\text{res},i}} \sum_{\ell \neq m} \frac{k_\ell}{k_m} \text{Re} \mathcal{I}_{K,i} - \frac{1}{X_{\text{res},i}^2} \frac{K(k_1^2 + k_2^2 + k_3^2)}{k_1 k_2 k_3} \text{Im} \mathcal{I}_{K,i} \right]. \quad (2.30)$$

Further we can evaluate $\mathcal{I}_{K,i}$ in the stationary phase approximation to get

$$\mathcal{I}_{K,i} = \frac{3b_i^\star \sqrt{2\pi}}{2} \left(\frac{\sqrt{2\epsilon_\star}}{f_i} \right)^{3/2} e^{i(\phi_K + c_i)/f_i} e^{-\frac{i\pi}{4}} e^{-i\left(\frac{\sqrt{2\epsilon_\star}}{f_i} - \frac{\sqrt{2\epsilon_\star}}{f_i} \ln \frac{\sqrt{2\epsilon_\star}}{f_i}\right)}. \quad (2.31)$$

Using $\phi_K = \phi_\star - \sqrt{2\epsilon_\star} \ln \frac{K}{k_\star}$, this gives us the shape of the non-Gaussianity as a sum of the resonant non-Gaussianity shapes in [18] (up to an overall phase):²

²Note that each term in the sum over i in eq. (2.32) actually has a different phase contribution which depends on the f_i , see eq. (2.31). However, we can absorb these terms into the c_i which we are free to choose, such that eq. (2.32) is correct.

$$\begin{aligned}
\frac{\mathcal{G}_{res}(k_1, k_2, k_3)}{k_1 k_2 k_3} &= \sum_i \frac{3\sqrt{2\pi} b_i^*}{8} \left(\frac{\sqrt{2\epsilon_\star}}{f_i} \right)^{3/2} \left[\sin \left(\frac{\phi_K + c_i}{f_i} \right) - \frac{f_i}{\sqrt{2\epsilon_\star}} \sum_{\ell \neq m} \frac{k_\ell}{k_m} \cos \left(\frac{\phi_K + c_i}{f_i} \right) \right. \\
&\quad \left. - \left(\frac{f_i}{\sqrt{2\epsilon_\star}} \right)^2 \frac{K(k_1^2 + k_2^2 + k_3^2)}{k_1 k_2 k_3} \sin \left(\frac{\phi_K + c_i}{f_i} \right) \right] \\
&= \sum_i \frac{3\sqrt{2\pi} b_i^*}{8} \left(\frac{\sqrt{2\epsilon_\star}}{f_i} \right)^{3/2} \left[\sin \left(\frac{\sqrt{2\epsilon_\star}}{f_i} \ln \frac{K}{k_\star} + \frac{c_i}{f_i} \right) + \right. \\
&\quad \left. \frac{f_i}{\sqrt{2\epsilon_\star}} \sum_{\ell \neq m} \frac{k_\ell}{k_m} \cos \left(\frac{\sqrt{2\epsilon_\star}}{f_i} \ln \frac{K}{k_\star} + \frac{c_i}{f_i} \right) + \dots \right].
\end{aligned} \tag{2.32}$$

As a consistency check of our calculation, we have checked that this sum of resonant bispectra satisfies the squeezed limit consistency relation [8, 23]. The details are presented in Appendix A.3.

3 Equilateral features from summed resonant non-Gaussianity

We now want to discuss what kind of non-Gaussianities could have a sizable overlap with summed resonant non-Gaussianity. One could also turn this question around and ask what values the parameters b_i^* , f_i and c_i have to take in order to generate a degeneracy. From an effective field theory point of view it is perfectly fine to choose ad hoc values for these parameters as long as they are not in conflict with observations such as the power spectrum and fulfill certain consistency conditions. For instance, the frequencies should certainly not be super Planckian, i.e. $f_i < 1$, and monotonicity of the inflaton potential requires $b_i^* < 1$. A possible stringy origin of these parameters is discussed below in Section 4.

In [24, 25], bounds on oscillating features in the power spectrum were given for a quadratic and a linear inflaton potential respectively. The bound that was found in both works is

$$b_i^* f_i < \frac{10^{-5}}{\sqrt{2\epsilon_\star}}, \tag{3.33}$$

for a single modulation in the potential. We assume that this bound is valid for multiple modulating terms in the potential. Regarding the applicability of the bound eq. (3.33) to general potentials $V_0(\phi)$, see the discussion in Section 3.4.

Let us briefly discuss the general form of the bispectrum. Due to the appearance of the delta function in eq. (2.25), a momentum configuration is completely characterized by the absolute values of the three momenta k_1 , k_2 and k_3 . Furthermore, for a scale-invariant spectrum this reduces to two variables. Then, one usually considers the bispectrum as a function of the two rescaled momenta $x_2 = k_2/k_1$ and $x_3 = k_3/k_1$. A region that includes only inequivalent momentum configurations is given by $1 - x_2 \leq x_3 \leq x_2$.

Let us switch to the variables $x_{\pm} = x_2 \pm x_3$, with $1 < x_+ < 2$ and $0 < x_- < 1$. Note that the resonant non-Gaussianity eq. (2.32) is to first order in $f_i/\sqrt{2\epsilon_{\star}}$ only a function of x_+ and k_1 but not of x_- since

$$\sin\left(\frac{\sqrt{2\epsilon_{\star}}}{f_i} \ln \frac{K}{k_{\star}}\right) = \sin\left(\frac{\sqrt{2\epsilon_{\star}}}{f_i} (y + \ln k_1/k_{\star})\right), \quad (3.34)$$

having defined $y \equiv \ln(1 + x_+)$. Note that summed resonant non-Gaussianities are therefore not scale invariant, as is clear from the explicit dependence on k_1 in eq. (3.34). We will discuss the issue of scale dependence in Section 3.2.

Furthermore, eq. (3.34) implies that as far as other types of non-Gaussianities are concerned, we can only expect degeneracies with the summed resonant type if they are predominantly a function of x_+ . We will show in the next section, Section 3.1, that this is primarily the case for equilateral non-Gaussianity, typically arising in non-canonical models of inflation.

To measure the degree of degeneracy between different kinds of non-Gaussianities we follow [26]. The bispectrum can always be parametrized in the form $\frac{\mathcal{G}(k_1, k_2, k_3)}{k_1 k_2 k_3} \sim f_{NL} S(k_1, k_2, k_3)$, with ‘amplitude’ f_{NL} and shape function $S(k_1, k_2, k_3)$. In general, the shape refers to the dependence of $S(k_1, k_2, k_3)$ on the momentum ratios k_2/k_1 and k_3/k_1 when the overall momentum scale K is fixed, while the dependence of S on K when the momenta k_i are fixed gives the running of the bispectrum.

Now the cosine of two shapes is defined via the normalized scalar product

$$C(S, S') = \frac{F(S, S')}{\sqrt{F(S, S)F(S', S')}}, \quad (3.35)$$

where

$$F(S, S') = \int_{\mathcal{V}} \frac{d\mathcal{V}}{K} S(k_1, k_2, k_3) S'(k_1, k_2, k_3), \quad (3.36)$$

with

$$\frac{d\mathcal{V}}{K} = \frac{dk_1 dk_2 dk_3}{k_1 + k_2 + k_3} = \frac{1}{2} k dk d\alpha d\beta, \quad (3.37)$$

where we have switched to the variables k , α and β defined according to [26]

$$\begin{aligned} k &= \frac{1}{2}(k_1 + k_2 + k_3), & k_1 &= k(1 - \beta), \\ k_2 &= \frac{k}{2}(1 + \alpha + \beta), & k_3 &= \frac{k}{2}(1 - \alpha + \beta), \end{aligned} \quad (3.38)$$

with the integration boundaries

$$\alpha \in [-1 + \beta, 1 - \beta], \quad \beta \in [0, 1] \quad \text{and} \quad k \in [k_{min}, k_{max}]. \quad (3.39)$$

3.1 Equilateral and local shapes for $x_- \rightarrow 0$

Out of the many known types of non-Gaussianities,³ we discuss the following two representative types: The equilateral type

$$\frac{\mathcal{G}_{equil}(k_1, k_2, k_3)}{k_1 k_2 k_3} = f_{NL}^{equil} S_{equil}(k_1, k_2, k_3), \quad (3.40)$$

with

$$S_{equil}(k_1, k_2, k_3) = \frac{(k_1 + k_2 - k_3)(k_1 + k_3 - k_2)(k_3 + k_2 - k_1)}{k_1 k_2 k_3}, \quad (3.41)$$

is characteristic of non-canonical inflation.

The local type is given by

$$\frac{\mathcal{G}_{local}(k_1, k_2, k_3)}{k_1 k_2 k_3} = f_{NL}^{local} S_{local}(k_1, k_2, k_3), \quad (3.42)$$

with

$$S_{local}(k_1, k_2, k_3) = \frac{k_1^3 + k_2^3 + k_3^3}{k_1 k_2 k_3}, \quad (3.43)$$

and is dominant in multi-field inflation for instance.

In the limit $x_- \rightarrow 0$, the shape functions are given by

$$S_{equil}^{x_- \rightarrow 0}(x_+) = \frac{4(x_+ - 1)}{x_+^2}, \quad (3.44)$$

$$S_{local}^{x_- \rightarrow 0}(x_+) = \frac{4 + x_+^3}{x_+^2}. \quad (3.45)$$

Now, one can evaluate the cosine eq. (3.35) to find the overlap of S_{equil} and $S_{equil}^{x_- \rightarrow 0}(x_+)$:

$$C(S_{equil}, S_{equil}^{x_- \rightarrow 0}) = 0.93. \quad (3.46)$$

³For an overview see e.g. [9].

Hence, the equilateral shape is well approximated by its $x_- \rightarrow 0$ limit. For the local shape the overlap is much smaller. S_{local} diverges for squeezed momentum configurations which makes it necessary to regulate the integrals that enter the cosine eq. (3.35). We only give an upper bound $C(S_{local}, S_{local}^{x_- \rightarrow 0}) < 0.7$, which was obtained by cutting the integration boundaries eq. (3.39) as follows:

$$\beta \in [\Delta, 1 - \Delta] \quad \text{with} \quad \Delta = 5 \cdot 10^{-5}. \quad (3.47)$$

The different qualitative levels of agreement of the equilateral and local shape are also visualized in Figure 2. Notice that the shape function $S_{non-can}$ from non-canonical inflation shows a slightly more complicated momentum dependence than the equilateral shape function eq. (3.41), see [16]. We find $C(S_{non-can}, S_{non-can}^{x_- \rightarrow 0}) > 0.93$, so the approximation by the $x_- \rightarrow 0$ limit is even better in the case of $S_{non-can}$.

3.2 Scale invariance vs. scale dependence

Let us now discuss how well the scale-invariant equilateral shape can be approximated by a scale-dependent shape, such as summed resonant non-Gaussianity. It is obvious that the overlap cannot be made arbitrarily large; however, we will show in the following that the overlap can still be considerable.

Let $S_{equil}^{per}(y)$ be the periodic generalization of

$$S_{equil}^{x_- \rightarrow 0}(y) = 4 \frac{e^y - 2}{(e^y - 1)^2}, \quad y \in [\ln 2, \ln 3], \quad (3.48)$$

to $y \in \mathbb{R}$, i.e. $S_{equil}^{per}(y + \Delta y) = S_{equil}^{per}(y)$ with $\Delta y = \ln 3 - \ln 2$. This definition is motivated by the fact that we are going to Fourier expand on the interval $y \in [\ln 2, \ln 3]$ and this expansion will itself be periodic with period Δy .

Furthermore, let us consider the scale-dependent shape $S_{equil}^{per}(y + \ln k_1/k_\star) = S_{equil}^{per}(\ln K/k_\star)$ which is a shape that can be approximated by the shape of some combination of resonant non-Gaussianities, since they have the same functional dependence, see eq. (3.34). The different shapes in $(y, \ln k_1/k_\star)$ space are shown in Figure 3.

The calculation of the cosine eq. (3.35) between the shapes $S_{equil}^{per}(\ln 2k/k_\star)$ and $S_{equil}(\alpha, \beta)$

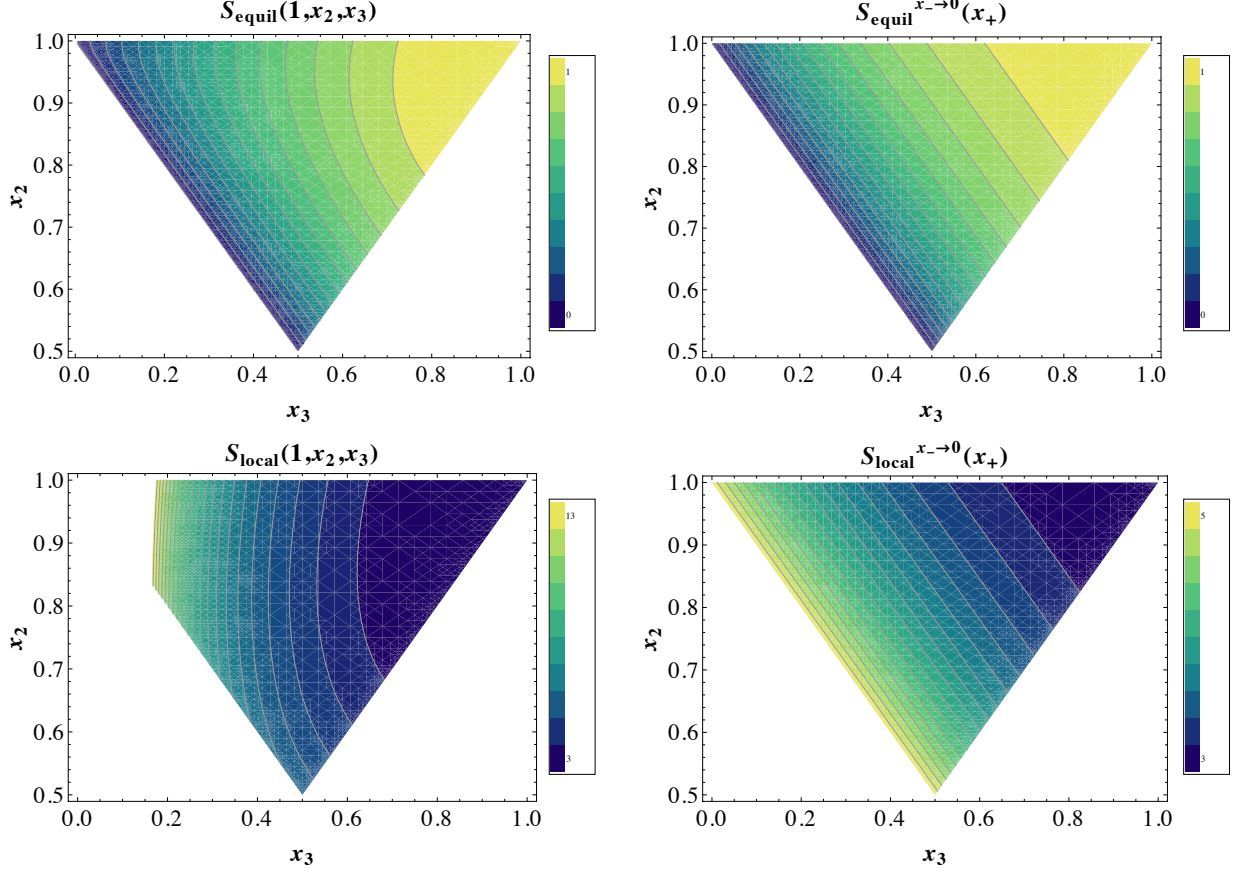


Figure 2: The equilateral and local shape functions (left) and their approximations $x_- \rightarrow 0$ (right).

simplifies due to the reduced dependencies on the integration variables k , α and β to

$$C(S_{equil}, S_{equil}^{per}) = \frac{\int d\alpha d\beta S_{equil}(\alpha, \beta)}{\left(2 \int d\alpha d\beta S_{equil}^2(\alpha, \beta)\right)^{1/2}} \frac{2 \int_{k_{min}}^{k_{max}} dk k S_{equil}^{per}(\ln 2k/k_*)}{\left([k_{max}^2 - k_{min}^2] \int_{k_{min}}^{k_{max}} dk k S_{equil}^{per}(\ln 2k/k_*)^2\right)^{1/2}} \quad (3.49)$$

We find that eq. (3.49) is to very good approximation independent of the values of k_{min} and k_{max} , if $k_{min} \ll k_*$ and $k_{max} \gg k_*$. This is due to the periodicity of $S_{equil}^{per}(\ln 2k/k_*)$, yielding the second fraction on the RHS of eq. (3.49) essentially independent of these values. For $k_{max} = 10^4 k_*$ and $k_{min} = 10^{-4} k_*$, we find

$$C(S_{equil}, S_{equil}^{per}) = 0.83. \quad (3.50)$$

Let us also mention that the overlap would be greater than 83%, if we had considered

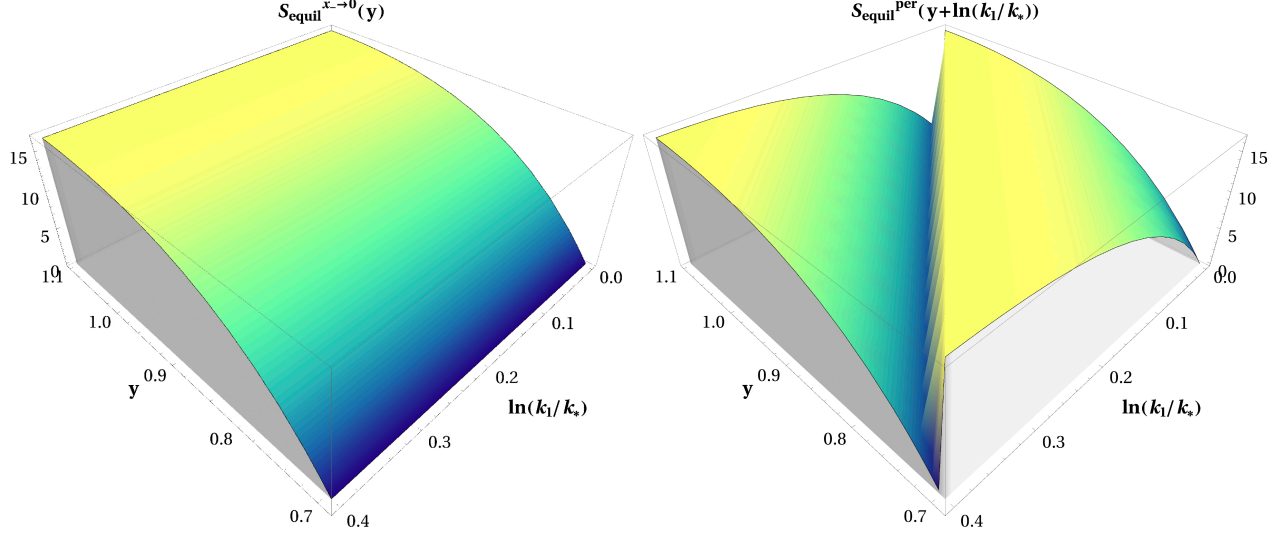


Figure 3: The scale-invariant shape $S_{equil}^{x \rightarrow 0}(y)$ (left) and scale-dependent shape $S_{equil}^{per}(y + \ln k_1/k_*)$ (right) for $y \in [\ln 2, \ln 3]$ and $\ln k_1/k_* \in [0, \Delta y]$

not the equilateral shape function eq. (3.41), but the exact shape function [16] that arises in non-canonical single field models of inflation. This is due to the fact that the overlap between $S^{x \rightarrow 0}$ and S is larger than for the equilateral shape function eq. (3.41), as discussed at the end of Section 3.1.

3.3 Fourier analysis

Having found that the equilateral shape function has a non-negligible degeneracy with the one dimensional function $S_{equil}^{per}(y + \ln k_1/k_*)$, the question of degeneracy with summed resonant non-Gaussianities can be expressed in terms of a Fourier analysis. Let us set $k_1 = k_*$ for conciseness in the remainder of this section; however the following analysis is valid for all values of k_1 .

Let us define

$$B_i = \frac{3\sqrt{2\pi}b_i^*}{8} \left(\frac{\sqrt{2\epsilon_*}}{f_i} \right)^{3/2} \quad \text{and} \quad F_i = \frac{\sqrt{2\epsilon_*}}{f_i}. \quad (3.51)$$

Then, summed resonant non-Gaussianity, eq. (2.32), can, to first order in $f_i/\sqrt{2\epsilon_*}$, be written

in the form⁴

$$\frac{\mathcal{G}_{res}}{k_1 k_2 k_3}(y) = \sum_{i=1}^{2N} B_i \sin(F_i y + C_i) = \sum_{i=1}^N B_i \cos(F_i y) + B_{N+i} \sin(F_i y) \quad (3.52)$$

where in the last equality of eq. (3.52), we have chosen the phases C_i such that there appears a sine and a cosine for each frequency F_i .

The Fourier expansion of a generic function $f(y)$ on the interval $y \in [a, a+b]$ is given as

$$f(y) \simeq \frac{u_0}{2} + \sum_i^N \left[u_i \cos\left(\frac{2\pi i}{b} y\right) + u_{N+i} \sin\left(\frac{2\pi i}{b} y\right) \right], \quad (3.53)$$

with

$$\begin{aligned} u_i &= \frac{2}{b} \int_a^{a+b} f(y) \cos\left(\frac{2\pi i}{b} y\right) dy \quad \text{for } i \geq 0, \\ u_{N+i} &= \frac{2}{b} \int_a^{a+b} f(y) \sin\left(\frac{2\pi i}{b} y\right) dy \quad \text{for } i \geq 1. \end{aligned} \quad (3.54)$$

Note that eq. (3.52) is of the Fourier form, eq. (3.53), if the constant term

$$u_0 = \frac{2}{b} \int_a^{a+b} f(y) dy, \quad (3.55)$$

is zero. Since $S_{equil}^{per}(y)$ is monotonically increasing and has a zero at $\ln 2$, the numbers a and b can always be chosen such that $u_0 = 0$. Note that we can always redefine $S_{equil}^{per}(y)$ on the (non-physical) interval $[\ln 2 - \xi, \ln 2]$ for $\xi \ll 1$ such that the integral eq. (3.55) has a large negative contribution on this interval, i.e.

$$u_0 \sim \int_a^{a+b} f(y) dy = \int_{\ln 2 - \xi}^{\ln 2} f(y) dy + \int_{\ln 2}^{\ln 3} f(y) dy, \quad (3.56)$$

where the first term on the RHS of eq. (3.56) is negative and has the same absolute value as the positive second term such that $u_0 = 0$. For instance one could use $f(y) = -|f_0|$ with $|f_0| \gg 1$ for $y \in [\ln 2 - \xi, \ln 2]$. Hence, $b = \ln 3 - \ln 2$.

In Figures 4 and 5, we show the Fourier expansion of $S_{equil}^{per}(y)$ for $N = 5$ and $N = 10$. The coefficients u_i are given in table 1 for $N = 5$.

⁴Due to the suppression by $f_i/\sqrt{2\epsilon_*} \ll 1$, the second order contribution to $\mathcal{G}_{res}/(k_1 k_2 k_3)$ induces a non-negligible correction only in the squeezed momentum configurations which we ignore in the following.

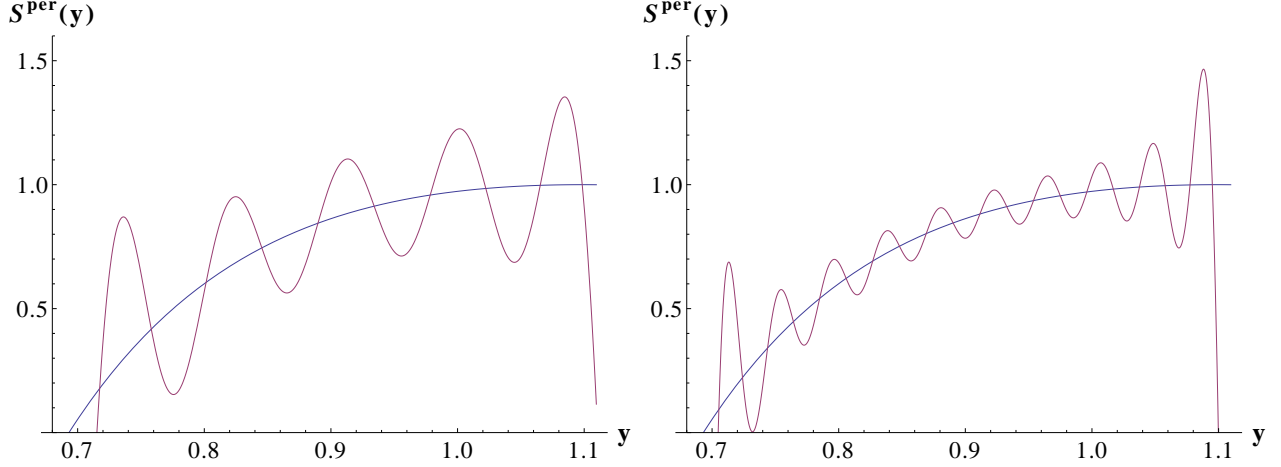


Figure 4: Fourier expansion of $S_{equil}^{per}(y)$ for $N = 5$ (left) and $N = 10$ (right).

u_1	u_2	u_3	u_4	u_5	u_6	u_7	u_8	u_9	u_{10}
0.41	0.01	-0.17	0.04	0.09	-0.20	0.24	-0.03	-0.12	0.05

Table 1: Fourier expansion coefficients u_i , eq. (3.54), of $S_{equil}^{per}(y)$ for $N = 5$.

Let us discuss how the Fourier expansion parameters a , b and u_i of eq. (3.54) relate to the parameters B_i and F_i of eq.'s (3.51) and (3.52). Clearly, if we want to impose

$$\mathcal{G}_{equil}^{per}(y) \simeq \mathcal{G}_{res}(y), \quad (3.57)$$

we have to set

$$B_i = f_{NL}^{equil} u_i \leq 0.41 f_{NL}^{equil} \quad \text{and} \quad F_i = \frac{2\pi i}{b} \simeq 15.5 i, \quad (3.58)$$

where we have used that the maximal u_i is $u_1 = 0.41$ for $N = 5$. In the derivation of the bispectrum in Section 2.3 we assumed $f_i \ll \sqrt{2\epsilon_\star}$, i.e. $F_i \gg 1$. We see from eq. (3.58) that this constraint is satisfied with growing i . Now, solving for b_i^\star and f_i in eq. (3.51) yields

$$b_i^\star = \frac{8}{3\sqrt{2\pi}} \frac{f_{NL}^{equil} u_i}{F_i^{3/2}} \leq \frac{f_{NL}^{equil}}{f_{NL}^{max}} \quad \text{and} \quad f_i = \frac{\sqrt{2\epsilon_\star}}{F_i} = 0.06 \frac{\sqrt{2\epsilon_\star}}{i}, \quad (3.59)$$

where we have once again used $\max\{u_i\} = 0.41$ and $f_{NL}^{max} \equiv 140$.

There are two things we can conclude from eq. (3.59). First of all, frequencies are necessarily sub-Planckian as is required for consistency. Second, we see that for $f_{NL}^{equil} > f_{NL}^{max}$ the monotonicity condition $b_i^\star < 1$ is violated. Hence $f_{NL}^{max} \equiv 140$ is the maximal equilateral

non-Gaussianity that could be matched by summed resonant non-Gaussianity if there were no constraints from the power spectrum. The f_{NL} that is consistent with the data is much smaller than 140, as we will discuss in the following section.

3.4 Constraints from the power spectrum

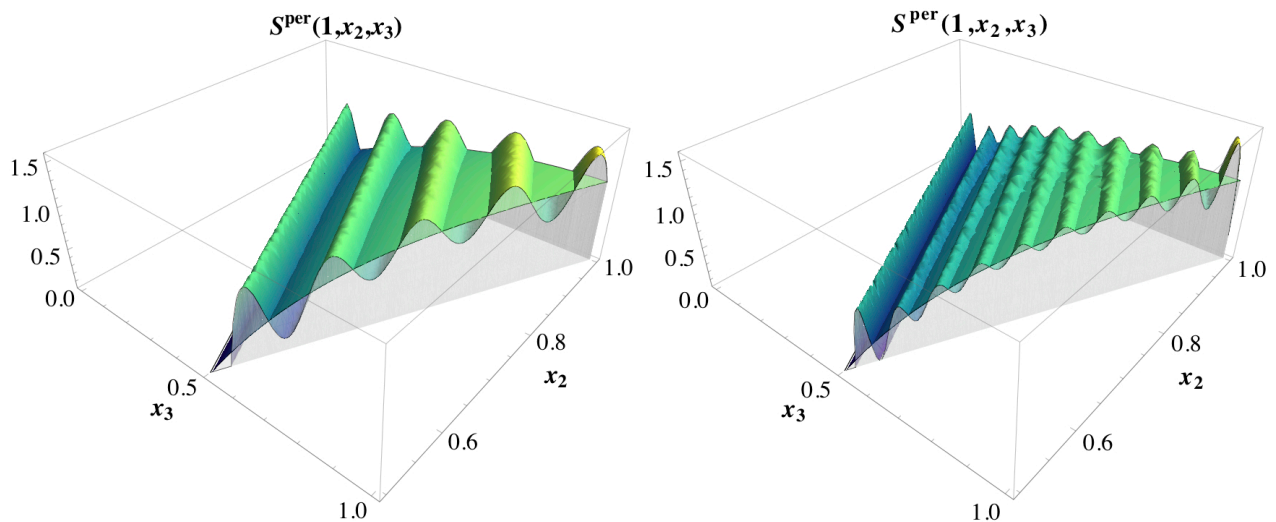


Figure 5: Fourier expansion of $S_{equil}^{per}(y)$ for $N = 5$ (left) and $N = 10$ (right). We have set $k_1 = k_*$.

Let us discuss the constraint on the product $b_i^* f_i$ from the power spectrum, i.e. eq. (3.33). Using eq. (3.59), we find

$$b_i^* f_i < \frac{8}{3\sqrt{2\pi}} \frac{\sqrt{2\epsilon_*} f_{NL}^{equil} u_i}{F_i^{5/2}} < 5 \cdot 10^{-4} f_{NL}^{equil} \sqrt{2\epsilon_*}, \quad (3.60)$$

where the product on the far right is the maximum value of $b_i^* f_i$. Since the maximum value of $b_i^* f_i$ has to be smaller than the upper bound given in eq. (3.33) this implies

$$\epsilon_* f_{NL}^{equil} < 10^{-2}. \quad (3.61)$$

Given the values for b_i^* and f_i obtained for the Fourier expansion, we can check the behaviour of eq. (2.22). We find that to leading order it behaves like a single oscillation with f_i given by the lowest frequency f_1 . This means that the two-point function bounds found for the singly modulated potential will still apply in our case.

Note, that eq. (3.33) was derived for large-field models (linear axion monodromy inflation [25] and a quadratic potential [24]), and thus should not have validity for, say, $\epsilon_\star \lesssim 10^{-3}$. Therefore, from this result we can argue at most for a resonantly generated $f_{NL}^{equil} \lesssim \mathcal{O}(1)$. More generally, an analysis of inflation in presence of periodically broken shift symmetry which generates resonant non-Gaussianity showed [20] that the structure of the resonant N -point functions displays a hierarchical suppression with increasing N . Hence, the non-observation of oscillating contributions to the 2-point function places tight constraints on the maximum amount of resonantly generated non-Gaussianity to no greater than the PLANCK sensitivity, that is $f_{NL} \sim \mathcal{O}(5)$.

4 Possible avenues for embedding: Axion monodromy inflation

We have seen that a potential modulated by a sum of small oscillating terms as given in eq. (2.3) :

$$V(\phi) = V_0(\phi) + \sum_i A_i \cos\left(\frac{\phi + c_i}{f_i}\right), \quad (4.62)$$

will lead to a bispectrum given by a sum of resonant bispectrum terms eq. (2.32). This follows as long as $f_i \ll \sqrt{2\epsilon_\star}$ for each i . Furthermore, a Fourier series of these resonant bispectra leads to a summed resonant bispectrum which closely approximates the equilateral bispectrum. Taking this series amounts to a choice of parameters $\{f_i, c_i\}$, consistent with the bounds above, for a suitable number of terms in the bispectrum or equivalently in the modulated potential eq. (2.3).

It remains to ask how such a sum of sinusoidal corrections to the potential could arise. In this section we argue for one possible source for these corrections in the context of axion monodromy inflation [27, 25], where the resonant non-Gaussianity shape was first observed [18]. The oscillatory correction considered in these papers arises from nonperturbative effects such as instantons, and is quite general in large field models. We show that these effects can also give rise to a sum of periodic potentials such as that in eq. (4.62) and studied here.

4.1 Axion monodromy inflation

Up to this point we have taken the potential eq. (4.62) as the starting point, in keeping with a low-energy effective field theory point of view of inflation in which we are agnostic as

to the UV origins of the theory. One possible avenue for a realization of multiple resonant non-Gaussianity is given by axion monodromy inflation [27, 25]. Axion monodromy inflation is a string theoretic realization of large field inflation, where the approximate shift symmetry of the axion protects the required flatness of the potential from higher-dimensional operators which become relevant when the field range is super-Planckian. Axions in string theory arise either as the Hodge duals of 2-form fields in 4 dimensions or as zero modes of p-forms on the compact manifold [28], and inherit their periodicity from these fields. Normalization of the axion kinetic term results in a coupling to other fields inversely proportional to f_a , the axion decay constant. In general string theoretic constructions this decay constant is larger than the range allowed by cosmological bounds (to avoid excessive production of axionic dark matter which could overclose the universe) [28]. However, despite this, the field range spanned within a single period is generally sub-Planckian in string theory [29].

This limitation (for the purposes of large field inflation) is circumvented in axion monodromy inflation, where a large field range for the inflaton axion field $c = 2\pi \int_{\Sigma} C_2$ is possible for instance when the two-cycle Σ on which it is supported is wrapped by a space-filling 5-brane which breaks the shift symmetry $c \rightarrow c + 2\pi$ in the potential, and gives rise to monodromy in the potential energy: the potential is no longer a periodic function of the axion, and increases without bound as the axion VEV increases. This follows from the DBI action, which gives a potential

$$V_0(c) \sim \sqrt{\ell^4 + c^2}, \quad (4.63)$$

where $\ell\sqrt{\alpha'}$ is the size of the wrapped cycle. Generalizations of this effect arise generically in the presence of fluxes.

The initial axion VEV is large, leading to a linear unmodulated potential $V_0(\phi)$, which means that $\frac{1}{\sqrt{2\epsilon_*}} = \phi_*$, the value of the inflaton at horizon exit, and $b^* = b$ is independent of ϕ_* . Inflation ends when at small values of the axion VEV, $V_0(\phi)$ is no longer linear, the axion oscillates about its minimum, and the inflationary energy is diverted into any string modes coupled to the axion.

Many of the consistency conditions (such as e.g. the absence of large back reaction from the axion-induced 3-brane charge building up on the 5-branes) for axion monodromy inflation were already discussed in [27, 25]. Some remaining questions may consist of e.g. finding an explicit realization on a concrete compact 3-fold in type IIB (despite the generality of the mechanism which in the context of fluxes [30] should lead to many explicit avenues for

construction), or details of the 3-brane charge back reaction and the mass scale of some of the KK modes in particular setups. For a more detailed discussion of some of these and related points see e.g. [25, 31].

4.2 Periodic corrections to the potential

In [27, 25], working in an $O3/O7$ orientifold of Type IIB, the axion found to be suitable for this model of inflation was that arising from the zero mode of the RR two-form C_2 on a 2-cycle Σ in the internal manifold, denoted $c = 2\pi \int_{\Sigma} C_2$. The cohomology groups $H^{(p,q)}$ split into two eigenstates under the action of the orientifold in a flux compactification [32] (see also [33])

$$H^{(p,q)} = H_+^{(p,q)} \oplus H_-^{(p,q)}, \quad (4.64)$$

where the \pm subscripts refer to even/odd behaviour under the orientifold action. Because C_2 is odd under the orientifold projection, only the mode coming from wrapping a 2-cycle which is also odd under the orientifold will survive, i.e. we take $\Sigma \in H_2^-$. Such an odd cycle can be written as $v^- = v^1 - v^2$, where v^1 and v^2 are two-cycles in the CY which are mapped into each other by the orientifold action. The $+$ combination is then obviously an even two-cycle.

Generally, one expects non-perturbative effects to lead to small periodic corrections to the potential for the axion, leading to the modulated potential which gives rise to resonant-type non-Gaussianity [25, 18]. These corrections arise from instanton corrections to the action, specifically worldsheet instanton contributions to the Kähler potential. Instantons arise when the $(p+1)$ -dimensional world volume of a Euclidean p -brane is wrapped over a $(p+1)$ cycle in the internal manifold. Their corrections are exponentially suppressed by the size of the wrapped cycles. Space-time instantons, which correct the superpotential, are localized in R^4 , while worldsheet instantons, which correct the Kähler potential, are localized in 2 dimensions. These are due to Euclidean D1-branes and their $SL(2, \mathbb{Z})$ images (i.e. F- and D-strings as well as their bound states, (p, q) strings) wrapping 2-cycles v^+ threaded by the axionic C_2 forms. In [25] this correction was taken to be of the form

$$\mathcal{K} = -2 \log[\mathcal{V}_E + e^{-2\pi v^+/\sqrt{g_s}} \cos(c)], \quad (4.65)$$

for the c axion, where \mathcal{V}_E is the Einstein-frame volume of the CY manifold, and v^+ is the even cycle dual to that wrapped by C_2 and on which the instanton is supported. τ is the

axion-dilaton $\tau = C_0 + ie^{-\phi}$ and the G^a are complex scalar fields composed of the RR and NSNS axionic fields c and b as [34]

$$G^a = c^a - \tau b^a, \quad (4.66)$$

where a runs over the number of axions.

The expression in eq. (4.65) captures the important sinusoidal features of the instanton corrections to the Kähler potential. The detailed form of one such series of leading (in the sense of exponentially unsuppressed) corrections in the b -axions was found in [35]. It is given by

$$g(\tau, \bar{\tau}, G^a, \bar{G}^a) = -\frac{1}{4} \sum_{\beta} n_{\beta} \sum_{(m,n)} \frac{(\tau - \bar{\tau})^{3/2}}{|n + m\tau|^3} \cos((n + \tau m) \frac{K_a(G^a - \bar{G}^a)}{\tau - \bar{\tau}} - mk_a G^a), \quad (4.67)$$

where β are curves in the negative eigenspace H_2^- spanned by the basis $\{\omega^a\}$, the n_{β} are the Gopakumar-Vafa invariants for these curves and $K_a = \int_{\beta} \omega_a$. The sum over (m, n) corresponds to a summation over $SL(2, \mathbb{Z})$ images of worldsheet instantons, which arise from F-strings wrapping 2-cycles. This ensures inclusion of Euclidean D1-branes and bound states, (p, q) -strings, in the instanton corrections. Generically then, at large $|b_i|$ this correction is a function of all the b_i -axions present.

On general grounds then (see also the short discussion in [35] on p. 9) we expect the presence of a qualitatively analogous series of corrections which depend on the c^i -axions at large $|c^i|$ and which are exponentially suppressed in the volumes of the curves in H_2^+ . As these terms will again contain a summation over all images under $SL(2, \mathbb{Z})$, the individual terms will not be $\sim \cos(c^i)$ for a given c^i but instead should have a dependence $\sim \cos(\alpha(n, m, \tau)_i c^i)$.

We may thus proceed by assuming the corrected Kähler potential to be of the form

$$\mathcal{K} = -2 \log \left\{ \mathcal{V}_E - e^{-\frac{2\pi v_+}{\sqrt{g_s}}} \sum_{n,m} A_{n,m}(\tau, \bar{\tau}) \cos[\alpha(n, m, \tau, \bar{\tau})c - \psi(m, n, \tau, \bar{\tau}, b)] \right\}, \quad (4.68)$$

for the case of one c -axion. The moduli potential often stabilizes b due to its appearance in the Kähler potential already at the leading level [34, 27]. Note, that this fixes b not necessarily at zero which can give rise to a non-zero phase $\psi(m, n, \tau, \bar{\tau}, b)$.

This correction to the Kähler potential results in a proportional correction to the scalar potential of the theory, as well as corrections to the kinetic terms of τ and G^a which are

exponentially suppressed in the curve volume v_+ . Hence, normalization of the axion kinetic term, as e.g. in (5.19) of [25], will yield a correction to the potential of the form of eq. (4.62)

$$\delta V \sim e^{-2\pi v_+/\sqrt{g_s}} \sum_i \tilde{A}_i \cos(\tilde{\alpha}_i \frac{\phi}{f} - \psi_i), \quad (4.69)$$

where we renamed the corresponding coefficients in eq. (4.68) as effective parameters $\tilde{A}_i, \tilde{\alpha}_i$.

For a very crude estimate of the number of terms one may expect to contribute to δV we note that already values of n, m smaller than 3 give $\mathcal{O}(10)$ possible terms in the $SL(2, \mathbb{Z})$ sum, each of which can have contributions from the sum over rational curves in H_2^- . This could easily give the $\mathcal{O}(20)$ terms required in the sum of resonant bispectra to reproduce an equilateral-type non-Gaussian signal.

Thus we have argued it to be plausible that a series of sinusoidal corrections with suitably varying frequencies may be achievable in connection with an axion monodromy generated potential, using the instanton corrections already considered at leading order in previous work. While the necessary and much deeper analysis of such constructions is left to the future, such a setup could form the starting point for a concrete model realizing the mechanism described in the preceding sections.

5 Conclusions

We have seen that a potential modulated by a sum of small oscillating terms will lead to a bispectrum given by a sum of resonant bispectrum terms which can sum to a bispectrum which closely approximates the equilateral bispectrum. This has potentially relevant implications for the aim of differentiating between inflationary models using data on non-Gaussianities. This potential differentiating power rests on (a) non-Gaussianity being one of the few possible consequences of inflation that is not generic, (b) different classes of inflationary theories giving rise to different dominant bispectrum shapes, and (c) new data allowing us to set bounds on the amplitude of different shapes in the bispectrum being expected in the next few years [36].

In broad terms, single scalar field inflationary models with canonical kinetic terms have non-Gaussian signal of order the slow-roll parameters, so that they would be ruled out by any measurable detection of non-Gaussianity. $P(X, \phi)$ theories (or more generally, Horn-deski theories) with non-canonical kinetic terms of a single inflaton field give rise to pre-

dominantly equilateral-type non-Gaussianity [16], while multi-field inflation gives rise to predominantly local-type non-Gaussianity and folded-type non-Gaussianity could indicate a non-Bunch Davies vacuum (For reviews see [9, 10] and references therein). Orthogonal-type non-Gaussianity, found in [37], can be the dominant type in multi-field DBI galileon inflation with an induced gravity term [38, 39]. Different contributions to the non-Gaussianity can also appear in combination, as in [40].

Here we have focussed on the possible degeneracy of equilateral-type non-Gaussianity, associated with single field theories with non-canonical kinetic terms, and summed resonant non-Gaussianity in a theory with canonical kinetic terms, arising from a potential modulated by a series of sinusoidal terms. It has been noted that within the class of non-canonical theories described by the Horndeski action, it may be difficult to differentiate between different models because some linear combination of the same shapes is always present [41, 42, 14]. The best way to differentiate between these models is therefore via conditions on the relative amplitudes of these terms which will differ for different models [43]. Despite this, one might hope to be able to identify the presence of non-canonical kinetic terms, via the dependence of the equilateral signal on the reduced sound speed $f_{NL}^{eq} \sim \frac{1}{c_s^2}$.

However, there is considerable degeneracy between observables even at this level. A dominantly equilateral signal has been found in multi-field models such as trapped inflation [44], where there is no reduction in the speed of sound. In addition, it was shown recently that f_{NL}^{eq} may not depend explicitly on c_s^2 in a weakly coupled extension of the EFT of inflation with reduced sound speed, depending instead on the Hubble scale and the UV cutoff of the theory [45]. Our result adds to the existing degeneracies as we have shown that even differentiating between canonical single field inflation models, with features in the potential, and non-canonical single field models may be more difficult than initially anticipated. However, differentiation is possible due to the tight constraints on the maximum amount of resonantly generated equilateral-type non-Gaussianity which arise from the non-observation of oscillating contributions to the 2-point function, and the structure of the resonant N -point functions. In particular, this implies that detection of equilateral non-Gaussianity at a level greater than the PLANCK sensitivity of $f_{NL} \sim \mathcal{O}(5)$ will rule out a resonant origin [20]. In case a detection of non-Gaussianity eludes PLANCK, future 21cm observations may provide a detection capability of $f_{NL} \gtrsim \mathcal{O}(0.01)$ down to the slow-roll level itself due to the extremely large number of modes ($\mathcal{O}(10^{16})$) available [21, 22]. For

equilateral-type non-Gaussianities of this magnitude there is then a degeneracy between a possible non-canonical and a resonant origin.

Let us note in closing that we have consciously kept most of discussion at the level of effective field theory. Beyond the very general considerations laid out in Section 4 any concrete analysis of embedding quasi-equilateral non-Gaussianity from multiple resonant oscillatory contributions into string theory e.g. along the lines of axion monodromy has to await much fuller treatment in the future.

Acknowledgments

The authors are grateful to E. Pajer for for several crucial and illuminating comments. The work of A.W. was supported by the Impuls und Vernetzungsfond of the Helmholtz Association of German Research Centres under grant HZ-NG-603, and German Science Foundation (DFG) within the Collaborative Research Center 676 “Particles, Strings and the Early Universe”. R.G. is grateful for support by the European Research Council via the Starting Grant numbered 256994. R.G. was also supported during the initial stages of this work by an SFB fellowship within the Collaborative Research Center 676 “Particles, Strings and the Early Universe” and would like to thank the theory groups at DESY and the University of Hamburg for their hospitality at this time.

A Computing the mode functions and bispectrum

A.1 The mode functions

We work in the comoving or unitary gauge where the inflationary perturbations are set to zero and spatial perturbations of the metric are written

$$\delta g_{ij}(\mathbf{x}, t) = 2a(t)^2 \mathcal{R}(\mathbf{x}, t) \delta_{ij}. \quad (\text{A.70})$$

The Fourier transform $\mathcal{R}(k, t)$ can be written in terms of modes

$$\mathcal{R}(\mathbf{k}, t) = \mathcal{R}_k(t) a(\mathbf{k}) + \mathcal{R}_k^*(t) a^\dagger(-\mathbf{k}), \quad (\text{A.71})$$

for which the mode function $\mathcal{R}_k(t)$ satisfies the Mukhanov-Sasaki equation [46, 47], which for small ϵ is given by [48]

$$\frac{d^2 \mathcal{R}_k}{dx^2} - \frac{2(1 + 2\epsilon + \delta)}{x} \frac{d\mathcal{R}_k}{dx} + \mathcal{R}_k = 0, \quad (\text{A.72})$$

where $x = -k\tau$, for τ the conformal time. For $x \ll 1$, $\mathcal{R}_k(t)$ approaches a constant value $\mathcal{R}_k^{(0)}$ which is related to the primordial scalar power spectrum. In the slow-roll approximation the Mukhanov-Sasaki equation is easily solved to give the answer in (2.16) of [18]. But in our case, as in [18], $\dot{\delta}/H$ is large compared to ϵ and δ and this SR approximation is invalid. Neglecting ϵ_0 and δ_0 and to leading order in ϵ_* , the Mukhanov-Sasaki equation becomes

$$\frac{d^2 \mathcal{R}_k}{dx^2} - \frac{2(1 + \delta_1(x))}{x} \frac{d\mathcal{R}_k}{dx} + \mathcal{R}_k = 0. \quad (\text{A.73})$$

for

$$\delta_1 = -3 \sum_i b_i^* \sin \left(\frac{\phi_0 + c_i}{f_i} \right). \quad (\text{A.74})$$

As in the single modulation case, the effect of δ_1 is negligible at early times, for $x \gg x_{\text{res},i} \forall i$, and

$$\mathcal{R}_k(x) = \mathcal{R}_{k,0}^{(o)} i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(1)}(x), \quad (\text{A.75})$$

where $\mathcal{R}_{k,0}^{(o)}$ is the value of $\mathcal{R}_k(x)$ outside the horizon in the absence of modulations and $H_{3/2}^{(1)}(x)$ is a Hankel function. Similarly for $x \ll x_{\text{res},i}$, i.e. at sufficiently late times, the effect of δ_1 is again negligible and the solution must be of the form

$$\mathcal{R}_k(x) = \mathcal{R}_{k,0}^{(o)} \left[c_k^{(+)} i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(1)}(x) - c_k^{(-)} i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(2)}(x) \right]. \quad (\text{A.76})$$

As in [25], $c_{k,i}^{(+)} = 1 + \mathcal{O}((b_\star^i)^2)$ at late times, where $c_k^{(+)} = \sum_i c_{k,i}^{(+)} = 1 + g(x)$ in the notation of [25]. (One can see this by solving (3.31) of [25] for each i).

Thus we look for a solution of the form

$$\mathcal{R}_k(x) = \mathcal{R}_{k,0}^{(o)} \left[i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(1)}(x) - c_k^{(-)}(x) i \sqrt{\frac{\pi}{2}} x^{3/2} H_{3/2}^{(2)}(x) \right], \quad (\text{A.77})$$

where $c_k^{(-)}(x)$ vanishes at early times. Then the Mukhanov-Sasaki equation gives an equation governing the time evolution of $c_k^{(-)}(x)$ which is given by (for $f_i \ll \sqrt{2\epsilon_*}$ for all i)

$$\frac{d}{dx} \left[e^{-2ix} \frac{d}{dx} c_k^{(-)}(x) \right] = -2i \frac{\delta_1(x)}{x}. \quad (\text{A.78})$$

This can be written as

$$\frac{d}{dx} \left[e^{-2ix} \frac{d}{dx} c_k^{(-)}(x) \right] = \frac{6i}{x} \sum_i b_i^* \sin\left(\frac{\phi_k}{f_i} + \frac{c_i}{f_i} + \frac{\sqrt{2\epsilon_\star}}{f_i} \ln x\right), \quad (\text{A.79})$$

$$\frac{d}{dx} A(x) = \frac{6i}{x} \sum_i b_i^* \sin\left(\frac{\phi_k}{f_i} + \frac{c_i}{f_i} + \frac{\sqrt{2\epsilon_\star}}{f_i} \ln x\right). \quad (\text{A.80})$$

Let $A(x) = \sum_i A_i(x)$, so that

$$\sum_i \frac{d}{dx} A_i(x) = \frac{6i}{x} \sum_i b_i^* \sin\left(\frac{\phi_k}{f_i} + \frac{c_i}{f_i} + \frac{\sqrt{2\epsilon_\star}}{f_i} \ln x\right) \quad (\text{A.81})$$

$$A_i(x) = -\frac{6ib_i^* f_i}{\sqrt{2\epsilon_\star}} \cos\left(\frac{\phi_k}{f_i} + \frac{c_i}{f_i} + \frac{\sqrt{2\epsilon_\star}}{f_i} \ln x\right) \quad (\text{A.82})$$

is a solution. Then write

$$A(x) = \sum_i A_i(x) = e^{-2ix} \frac{d}{dx} c_k^{(-)}(x) = e^{-2ix} \frac{d}{dx} \sum_i c_{k,i}^{(-)}(x). \quad (\text{A.83})$$

Then we find

$$\frac{d}{dx} c_{k,i}^{(-)}(x) = -\frac{6ib_i^* f_i}{\sqrt{2\epsilon_\star}} e^{2ix} \cos\left(\frac{\phi_k}{f_i} + \frac{c_i}{f_i} + \frac{\sqrt{2\epsilon_\star}}{f_i} \ln x\right). \quad (\text{A.84})$$

To find the mode functions outside the horizon, we use the expansions of the Hankel functions for small arguments:

$$H_{3/2}^{(1)}(x) = \frac{4}{3\sqrt{\pi}} \left(\frac{x}{2}\right)^{3/2} - \frac{i}{2\sqrt{\pi}} \left(\frac{2}{x}\right)^{3/2} \quad (\text{A.85})$$

$$H_{3/2}^{(2)}(x) = \frac{4}{3\sqrt{\pi}} \left(\frac{x}{2}\right)^{3/2} + \frac{i}{2\sqrt{\pi}} \left(\frac{2}{x}\right)^{3/2}. \quad (\text{A.86})$$

We find

$$\begin{aligned} \mathcal{R}_k(x) = R_{k,0}^{(0)} & \left[1 + \sum_i 3b_i^* \sqrt{\frac{f_i \pi}{2\sqrt{2\epsilon_\star}}} \cos\left(\frac{\phi_k + c_i}{f_i}\right) \right. \\ & \left. - i \sum_i 3b_i^* \sqrt{\frac{f_i \pi}{2\sqrt{2\epsilon_\star}}} \sin\left(\frac{\phi_k + c_i}{f_i}\right) + \mathcal{O}(x^3) \right], \end{aligned} \quad (\text{A.87})$$

and

$$|\mathcal{R}_k^{(0)}|^2 = |\mathcal{R}_{k,0}^{(0)}|^2 \left[1 + \sum_i 3b_i^* \left(\frac{2f_i \pi}{\sqrt{2\epsilon_\star}}\right)^{1/2} \cos\left(\frac{\phi_k + c_i}{f_i}\right) \right]. \quad (\text{A.88})$$

A.2 Mode functions in the bispectrum

To see that we can use the unperturbed mode functions to calculate the bispectrum in Section 2.3, note first that at late times, for $x \ll 1$, the mode function in eq. (A.88) approaches a constant, given by $\mathcal{R}_{k,0}^{(o)}[1 + c_k^{(-)}]$. The contribution at early times, $x \gg 1$, within the horizon, is

$$\begin{aligned}\mathcal{R}_k(x) &= \mathcal{R}_{k,0}^{(0)} \left[-ixe^{ix} + ic_k^{(-)}(x)xe^{-ix} \right] \\ &= \mathcal{R}_{k,0} + \mathcal{R}_{k,1},\end{aligned}\tag{A.89}$$

where we have used the $x \gg 1$ limit of the Hankel functions [18]:

$$\begin{aligned}i\sqrt{\frac{\pi}{2}}x^{3/2}H_{3/2}^{(1)}(x) &\approx -ixe^{ix} \\ i\sqrt{\frac{\pi}{2}}x^{3/2}H_{3/2}^{(2)}(x) &\approx -ixe^{-ix}\end{aligned}\tag{A.90}$$

and $\mathcal{R}_{k,0}$ is the unperturbed part of the mode function. We find

$$\dot{\mathcal{R}}_{k,0} = \mathcal{R}_{k,0}^{(0)}xe^{ix}\tag{A.91}$$

$$\dot{\mathcal{R}}_{k,1} = \mathcal{R}_{k,0}^{(0)}xe^{-ix} \left[c_k^{(-)}(x) + i\frac{d}{dx}c_k^{(-)}(x) \right]\tag{A.92}$$

$$\Rightarrow \frac{|\dot{\mathcal{R}}_{k,1}|}{|\dot{\mathcal{R}}_{k,0}|} = \left| c_k^{(-)}(x) + i\frac{d}{dx}c_k^{(-)}(x) \right|.\tag{A.93}$$

By eq. (2.18) and eq. (A.84) we see that this ratio is small for $f_i \ll \sqrt{2\epsilon_\star}$ (which is when the non-Gaussianity is appreciable). Thus the correction to the time derivative of the mode function is also small, and we can use the unperturbed mode functions and time derivatives in evaluating the three-point function.

A.3 Squeezed limit consistency relation

In [18] the squeezed limit consistency relation found in [8, 23] was checked for resonant non-Gaussianity. It is straightforward to extend this to summed non-Gaussianities. The consistency condition can be phrased as [18]

$$\lim_{k_3 \rightarrow 0} \langle \mathcal{R}(\mathbf{k}_1, t) \mathcal{R}(\mathbf{k}_2, t) \mathcal{R}(\mathbf{k}_3, t) \rangle \simeq -(2\pi)^3 \delta^3(\mathbf{k}_1 + \mathbf{k}_2 + \mathbf{k}_3) |\mathcal{R}_{k_3}^{(0)}|^2 |\mathcal{R}_k^{(0)}|^2 \frac{d \ln \Delta_{\mathcal{R}}^2(k)}{d \ln k}, \tag{A.94}$$

where in the squeezed limit $k_1 \simeq k_2 = k \gg k_3$. Furthermore, $|\mathcal{R}_k^{(0)}|^2 = 2\pi^2 \Delta_{\mathcal{R}}^2/k^3$, with⁵

$$\Delta_{\mathcal{R}}^2(k) = \Delta_{\mathcal{R}}^2(k_*) \left(\frac{k}{k_*} \right)^{n_s-1} \left[1 + \sum_i 3b_i^* \sqrt{2\pi} \left(\frac{f_i}{\sqrt{2\epsilon_*}} \right)^{1/2} \cos \left(\frac{\sqrt{2\epsilon_*}}{f_i} \ln \frac{k}{k_*} \right) \right], \quad (\text{A.95})$$

such that

$$\frac{d \ln \Delta_{\mathcal{R}}^2(k)}{d \ln k} \simeq - \sum_i 3b_i^* \sqrt{2\pi} \left(\frac{\sqrt{2\epsilon_*}}{f_i} \right)^{1/2} \sin \left(\frac{\sqrt{2\epsilon_*}}{f_i} \ln \frac{k}{k_*} \right), \quad (\text{A.96})$$

neglecting slow roll corrections to $n_s = 1$. Now we can write

$$\begin{aligned} \lim_{k_3 \rightarrow 0} \langle \mathcal{R}(\mathbf{k}_1, t) \mathcal{R}(\mathbf{k}_2, t) \mathcal{R}(\mathbf{k}_3, t) \rangle &\simeq (2\pi)^7 \Delta_{\mathcal{R}}^4 \frac{1}{k_1 k_2 k_3} \delta^3(\mathbf{k}_1 + \mathbf{k}_2 + \mathbf{k}_3) \\ &\times \frac{2k}{k_3} \sum_i \frac{3b_i^* \sqrt{2\pi}}{8} \left(\frac{\sqrt{2\epsilon_*}}{f_i} \right)^{1/2} \sin \left(\frac{\sqrt{2\epsilon_*}}{f_i} \ln \frac{k}{k_*} \right). \end{aligned} \quad (\text{A.97})$$

Comparing with eq. (2.25), to fulfill the consistency condition eq. (A.94), \mathcal{G}_{res} has to be given as

$$\frac{\mathcal{G}_{res}(k, k, k_3)}{k^2 k_3} \simeq \frac{2k}{k_3} \sum_i \frac{3b_i^* \sqrt{2\pi}}{8} \left(\frac{\sqrt{2\epsilon_*}}{f_i} \right)^{1/2} \sin \left(\frac{\sqrt{2\epsilon_*}}{f_i} \ln \frac{k}{k_*} \right), \quad (\text{A.98})$$

in the squeezed limit. This is indeed the case up to a phase, as can be seen by taking the limit $k_3 \rightarrow 0$ of eq. (2.32).

References

- [1] **WMAP Collaboration** Collaboration, E. Komatsu *et. al.*, *Seven-Year Wilkinson Microwave Anisotropy Probe (WMAP) Observations: Cosmological Interpretation*, *Astrophys.J.Suppl.* **192** (2011) 18, [arXiv:1001.4538].
- [2] J. Dunkley, R. Hlozek, J. Sievers, V. Acquaviva, P. Ade, *et. al.*, *The Atacama Cosmology Telescope: Cosmological Parameters from the 2008 Power Spectra*, *Astrophys.J.* **739** (2011) 52, [arXiv:1009.0866].
- [3] R. Keisler, C. Reichardt, K. Aird, B. Benson, L. Bleem, *et. al.*, *A Measurement of the Damping Tail of the Cosmic Microwave Background Power Spectrum with the South Pole Telescope*, *Astrophys.J.* **743** (2011) 28, [arXiv:1105.3182].

⁵For this discussion we set the phases $c_i = 0$.

- [4] **Supernova Search Team** Collaboration, A. G. Riess *et. al.*, *Observational evidence from supernovae for an accelerating universe and a cosmological constant*, *Astron.J.* **116** (1998) 1009–1038, [astro-ph/9805201].
- [5] **Supernova Cosmology Project** Collaboration, S. Perlmutter *et. al.*, *Measurements of Omega and Lambda from 42 high redshift supernovae*, *Astrophys.J.* **517** (1999) 565–586, [astro-ph/9812133]. The Supernova Cosmology Project.
- [6] **SDSS Collaboration** Collaboration, W. J. Percival *et. al.*, *Baryon Acoustic Oscillations in the Sloan Digital Sky Survey Data Release 7 Galaxy Sample*, *Mon.Not.Roy.Astron.Soc.* **401** (2010) 2148–2168, [arXiv:0907.1660].
- [7] A. G. Riess, L. Macri, S. Casertano, H. Lampeitl, H. C. Ferguson, *et. al.*, *A 3Hubble Space Telescope and Wide Field Camera 3*, *Astrophys.J.* **730** (2011) 119, [arXiv:1103.2976].
- [8] J. M. Maldacena, *Non-Gaussian features of primordial fluctuations in single field inflationary models*, *JHEP* **0305** (2003) 013, [astro-ph/0210603].
- [9] X. Chen, *Primordial Non-Gaussianities from Inflation Models*, *Adv.Astron.* **2010** (2010) 638979, [arXiv:1002.1416].
- [10] K. Koyama, *Non-Gaussianity of quantum fields during inflation*, *Class.Quant.Grav.* **27** (2010) 124001, [arXiv:1002.0600].
- [11] E. Silverstein and D. Tong, *Scalar speed limits and cosmology: Acceleration from D-cceleration*, *Phys.Rev.* **D70** (2004) 103505, [hep-th/0310221].
- [12] M. Alishahiha, E. Silverstein, and D. Tong, *DBI in the sky*, *Phys.Rev.* **D70** (2004) 123505, [hep-th/0404084].
- [13] P. Franche, R. Gwyn, B. Underwood, and A. Wissanji, *Attractive Lagrangians for Non-Canonical Inflation*, *Phys.Rev.* **D81** (2010) 123526, [arXiv:0912.1857].
- [14] S. Renaux-Petel, *On the redundancy of operators and the bispectrum in the most general second-order scalar-tensor theory*, *JCAP* **1202** (2012) 020, [arXiv:1107.5020].
- [15] R. H. Ribeiro, *Inflationary signatures of single-field models beyond slow-roll*, *JCAP* **1205** (2012) 037, [arXiv:1202.4453].

- [16] X. Chen, M.-x. Huang, S. Kachru, and G. Shiu, *Observational signatures and non-Gaussianities of general single field inflation*, *JCAP* **0701** (2007) 002, [[hep-th/0605045](#)].
- [17] X. Chen, R. Easther, and E. A. Lim, *Generation and Characterization of Large Non-Gaussianities in Single Field Inflation*, *JCAP* **0804** (2008) 010, [[arXiv:0801.3295](#)].
- [18] R. Flauger and E. Pajer, *Resonant Non-Gaussianity*, *JCAP* **1101** (2011) 017, [[arXiv:1002.0833](#)].
- [19] L. Leblond and E. Pajer, *Resonant Trispectrum and a Dozen More Primordial N-point functions*, *JCAP* **1101** (2011) 035, [[arXiv:1010.4565](#)].
- [20] S. R. Behbahani, A. Dymarsky, M. Mirbabayi, and L. Senatore, *(Small) Resonant non-Gaussianities: Signatures of a Discrete Shift Symmetry in the Effective Field Theory of Inflation*, [arXiv:1111.3373](#).
- [21] A. Loeb and M. Zaldarriaga, *Measuring the small - scale power spectrum of cosmic density fluctuations through 21 cm tomography prior to the epoch of structure formation*, *Phys.Rev.Lett.* **92** (2004) 211301, [[astro-ph/0312134](#)].
- [22] A. Cooray, *21-cm Background Anisotropies Can Discern Primordial Non-Gaussianity*, *Phys.Rev.Lett.* **97** (2006) 261301, [[astro-ph/0610257](#)].
- [23] P. Creminelli and M. Zaldarriaga, *Single field consistency relation for the 3-point function*, *JCAP* **0410** (2004) 006, [[astro-ph/0407059](#)].
- [24] C. Pahud, M. Kamionkowski, and A. R. Liddle, *Oscillations in the inflaton potential?*, *Phys.Rev.* **D79** (2009) 083503, [[arXiv:0807.0322](#)].
- [25] R. Flauger, L. McAllister, E. Pajer, A. Westphal, and G. Xu, *Oscillations in the CMB from Axion Monodromy Inflation*, *JCAP* **1006** (2010) 009, [[arXiv:0907.2916](#)].
- [26] J. Fergusson and E. Shellard, *The shape of primordial non-Gaussianity and the CMB bispectrum*, *Phys.Rev.* **D80** (2009) 043510, [[arXiv:0812.3413](#)].
- [27] L. McAllister, E. Silverstein, and A. Westphal, *Gravity Waves and Linear Inflation from Axion Monodromy*, *Phys.Rev.* **D82** (2010) 046003, [[arXiv:0808.0706](#)].

- [28] P. Svrcek and E. Witten, *Axions In String Theory*, *JHEP* **0606** (2006) 051, [[hep-th/0605206](#)].
- [29] T. Banks, M. Dine, P. J. Fox, and E. Gorbatov, *On the possibility of large axion decay constants*, *JCAP* **0306** (2003) 001, [[hep-th/0303252](#)].
- [30] X. Dong, B. Horn, E. Silverstein, and A. Westphal, *Simple exercises to flatten your potential*, *Phys.Rev.* **D84** (2011) 026011, [[arXiv:1011.4521](#)].
- [31] N. Barnaby, E. Pajer, and M. Peloso, *Gauge Field Production in Axion Inflation: Consequences for Monodromy, non-Gaussianity in the CMB, and Gravitational Waves at Interferometers*, *Phys.Rev.* **D85** (2012) 023525, [[arXiv:1110.3327](#)].
- [32] S. B. Giddings, S. Kachru, and J. Polchinski, *Hierarchies from fluxes in string compactifications*, *Phys.Rev.* **D66** (2002) 106006, [[hep-th/0105097](#)].
- [33] K. Dasgupta, G. Rajesh, and S. Sethi, *M theory, orientifolds and G - flux*, *JHEP* **9908** (1999) 023, [[hep-th/9908088](#)].
- [34] T. W. Grimm and J. Louis, *The Effective action of $N = 1$ Calabi-Yau orientifolds*, *Nucl.Phys.* **B699** (2004) 387–426, [[hep-th/0403067](#)].
- [35] T. W. Grimm, *Non-Perturbative Corrections and Modularity in $N=1$ Type IIB Compactifications*, *JHEP* **0710** (2007) 004, [[arXiv:0705.3253](#)].
- [36] **Planck Collaboration** Collaboration, P. Ade *et. al.*, *Planck Early Results. I. The Planck mission*, *Astron.Astrophys.* **536** (2011) 16464, [[arXiv:1101.2022](#)].
- [37] L. Senatore, K. M. Smith, and M. Zaldarriaga, *Non-Gaussianities in Single Field Inflation and their Optimal Limits from the WMAP 5-year Data*, *JCAP* **1001** (2010) 028, [[arXiv:0905.3746](#)].
- [38] S. Renaux-Petel, *Orthogonal non-Gaussianities from Dirac-Born-Infeld Galileon inflation*, *Class.Quant.Grav.* **28** (2011) 182001, [[arXiv:1105.6366](#)].
- [39] S. Renaux-Petel, S. Mizuno, and K. Koyama, *Primordial fluctuations and non-Gaussianities from multifield DBI Galileon inflation*, *JCAP* **1111** (2011) 042, [[arXiv:1108.0305](#)].

- [40] D. Langlois, S. Renaux-Petel, D. A. Steer, and T. Tanaka, *Primordial perturbations and non-Gaussianities in DBI and general multi-field inflation*, *Phys.Rev.* **D78** (2008) 063523, [[arXiv:0806.0336](#)].
- [41] X. Gao and D. A. Steer, *Inflation and primordial non-Gaussianities of 'generalized Galileons'*, *JCAP* **1112** (2011) 019, [[arXiv:1107.2642](#)].
- [42] A. De Felice and S. Tsujikawa, *Inflationary non-Gaussianities in the most general second-order scalar-tensor theories*, *Phys.Rev.* **D84** (2011) 083504, [[arXiv:1107.3917](#)].
- [43] R. H. Ribeiro and D. Seery, *Decoding the bispectrum of single-field inflation*, *JCAP* **1110** (2011) 027, [[arXiv:1108.3839](#)].
- [44] D. Green, B. Horn, L. Senatore, and E. Silverstein, *Trapped Inflation*, *Phys.Rev.* **D80** (2009) 063533, [[arXiv:0902.1006](#)].
- [45] R. Gwyn, G. A. Palma, M. Sakellariadou, and S. Sypsas, *Effective field theory of weakly coupled inflationary models*, [arXiv:1210.3020](#).
- [46] V. F. Mukhanov, *Gravitational Instability of the Universe Filled with a Scalar Field*, *JETP Lett.* **41** (1985) 493–496.
- [47] M. Sasaki, *Large Scale Quantum Fluctuations in the Inflationary Universe*, *Prog.Theor.Phys.* **76** (1986) 1036.
- [48] S. Weinberg, *Cosmology*, .